

Dilational fault slip and pit chain formation on Mars

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ABSTRACT

Pit crater chains that are parallel or collinear with normal faults are common on Mars, but only rarely observed on Earth. We explore the origin of pit crater chains using recent high-resolution imagery from Mars, laboratory simulation of pit chains, investigation of recent pit chains formed in response to seismic fault slip in Iceland, and assessment of rock failure criteria and the Martian crustal stress field to develop a deeper understanding of the relationship between pit chain formation and fault activity. Based on these analyses, we conclude that pit chains form in response to dilational fault slip. Because gravitational acceleration on Earth (9.81 m/s^2) is higher than on Mars (3.72 m/s^2), stress within Earth is greater than that of Mars for any given depth. Consequently, steep fault segments are likely to extend to about 5 km depth on Mars compared with about 2 km depth on Earth. Dilation of these steep segments associated with fault slip on Mars could result in large volume increase in the uppermost crust, influencing groundwater flow and discharge and mineralization. Based on crosscutting relationships, pristine pit morphologies, and lack of evidence of sediment accumulation in the bottoms of pits, we interpret that some pit craters may be actively forming in response to dilational fault slip.

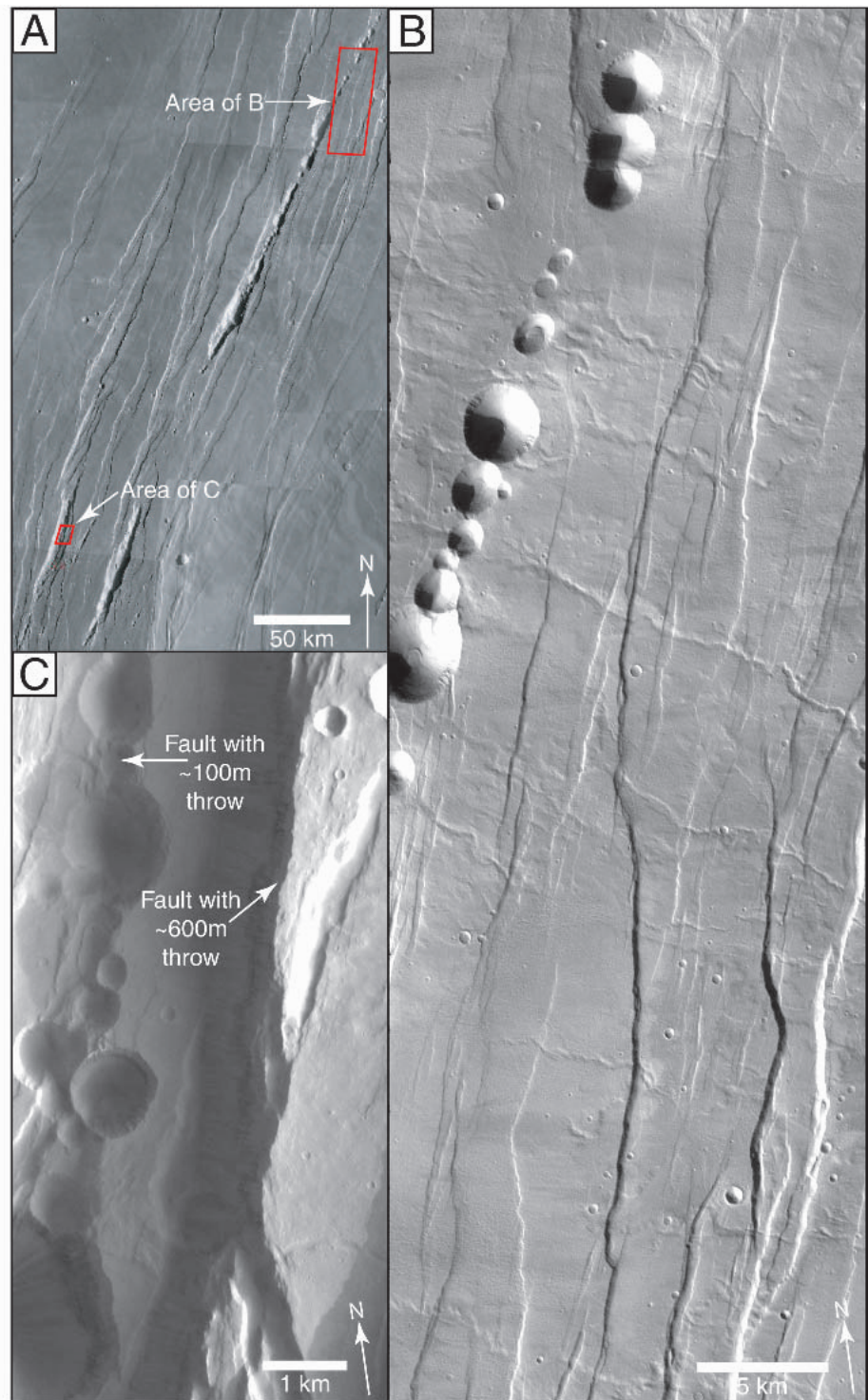


Figure 1. (A) Composite Viking imagery of normal fault scarps defining horsts and grabens in the southeast part of Alba Patera, Mars. Fault segmentation and en echelon arrangement of faults and grabens are common. Scale bar is 50 km. (B) THEMIS visible spectrum image shows detail of normal faults and pit chain (Phlegethon Catena) shown in A. Scale bar is 5 km. (C) Mars Orbital Camera image shows clear association of pit chains with normal fault scarps within graben on the southeast flank of Alba Patera. Scale bar is 1 km.

INTRODUCTION

The surface of Mars has been a source of debate since the earliest telescopic observations in the 1600s by Galileo, Huygens, and Cassini. Unprecedented detail in images and topographic maps generated by Mars missions within the past three decades allows identification of features carved by flowing water (Carr, 1996), accumulations of wind-blown sand and active dust devils (Malin and Edgett, 2001), and landforms generated by tectonic faulting (Schultz, 1985; Anderson et al., 2001; Wilkins and Schultz, 2003). Views of strata exposed in walls of canyons and pits demonstrate the presence of rock layers with different strengths (McEwen et al., 1999), dispelling the belief that the Martian surface has been so heavily damaged by impact cratering that it consists primarily of cohesionless regolith (Tanaka and Golombek, 1989).

Comparisons of geologic features on Mars with analogous features on Earth have been used extensively to understand Martian crustal processes. One distinctive feature of Martian topography is the occurrence of pit craters and pit chains (Tanaka and Golombek, 1989; Banerdt et al., 1992). Pit craters lack a raised crater rim or ejecta deposits, form alignments (chains), and are likely the result of collapse of loose surface material into a subsurface void (see Wyrick et al., 2004, and references therein). The close association of pit crater chains with faulting on Mars (Tanaka and Golombek, 1989; Wyrick et al., 2004) indicates that some Martian faults produce considerable subsurface void space. Collapse features seen on Earth are neither as widespread nor as large as those seen on Mars, nor is the mechanism of widespread dissolution of soluble rock, such as limestone or evaporites, observed on Earth likely on Mars. High-resolution imagery reveals that many pit craters have conical shapes. In addition, the walls of pit craters, troughs, grabens, and erosional canyons display a near-horizontal stratigraphy, composed of alternating layers with relatively low (gentle slopes) and relatively high (steep slopes) mechanical strength. The extreme size of pit craters on Mars, with diameters of individual pits reaching >4 km, and pit volumes >7 km³ (Wyrick et al., 2004), has remained an enigma.

Using recent high-resolution imagery from Mars, laboratory simulation of pit chains, investigation of recent pit chains formed in response to seismic fault slip in Iceland, and assessment of rock failure criteria and the Martian crustal stress field, we explore the origin of pit crater chains. We conclude that pit chains form in response to fault slip and dilation, consistent with the interpretation of active faulting on Mars (Anderson et al., 2001).

FAULT-LOCALIZED PIT CHAINS ON ALBA PATERA, MARS

Alba Patera is a large shield volcano, measuring 2700 km across, in the northern part of the Tharsis region on Mars (Cailleau et al., 2003). Normal faulting of Alba Patera is thought to have been active between the onset of the Noachian through the Early Amazonian—possibly more than a billion years of activity (Cailleau et al., 2003). Normal faults generally strike NNE-SSW through Alba Patera, but are tangential or locally circumferential to the volcano on the east, west, and northwest flanks of the volcano. Normal faults and associated horsts and grabens (fossae) locally host pit crater chains on

the eastern (Phlegethon Catenae, Archeron Catenae, Tractus Catenae) and western (Alba Catenae, Cyane Catenae, and Artynia Catenae) flanks of Alba Patera.

Images from Viking (e.g., Fig. 1A), Thermal Emission Imaging System (THEMIS; Fig. 1B), and Mars Orbital Camera (MOC; Fig. 1C) show the clear association of pit chains with normal faults, and the Phlegethon Catenae pit chain system displays the association of pits with a graben. Within a graben, pits and troughs tend to be localized along surface fault traces (Fig. 1C). The clear isolated pits in Fig. 1B are along a portion of Phlegethon Catenae where associated faults have relatively small displacements. Northeast and southwest of the area of Fig. 1B (see Fig. 1A) displacement increases laterally along the graben bounding faults and the pit chain has developed into a trough along the fault scarps. In the narrow-angle MOC image shown in Fig. 1C, pits are localized along a fault scarp with 100 m throw (based on Mars Orbiter Laser Altimeter, or MOLA, data). In the same figure, the fault with 600 m throw hosts a trough that we interpret to have developed from coalescence of a pit chain along the fault trace (Fig. 2A). The association of normal faults and pit chains (Fig. 2) implies that pit chains are surface indicators of normal faults which in turn influence

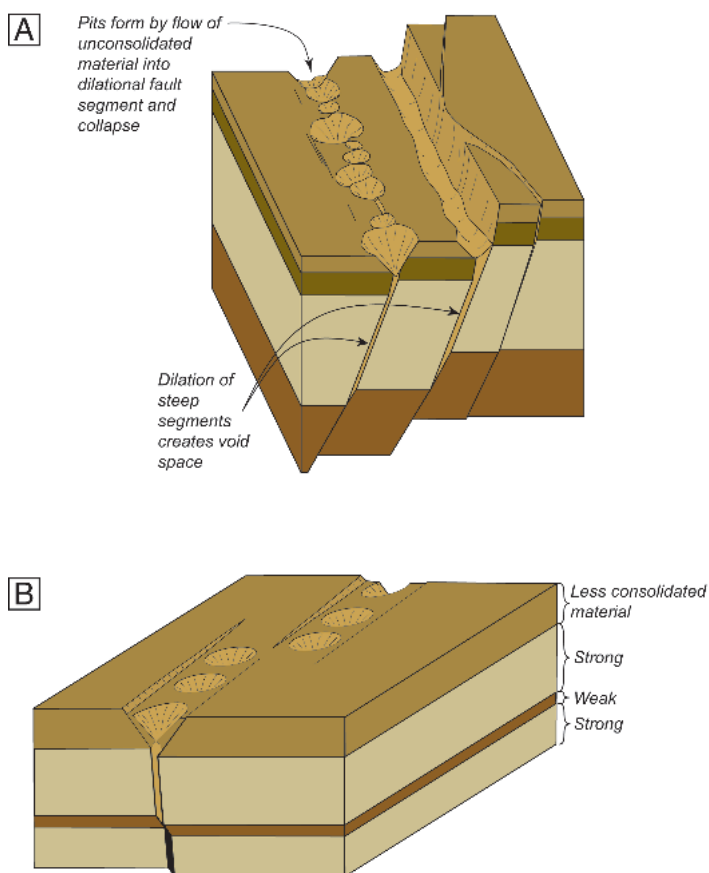


Figure 2. Schematic illustrations of faulting on Mars that display steep dilated faults and their potential relationship to pit chain development. (A) Pit chain and trough localized along the surface traces of dilational faults (compare with Fig. 1C). (B) Pit chain in surface graben in unconsolidated material above a dilational fault at depth (compare with Fig. 1A and B).

hydrothermal circulation, groundwater hydrology, and seismicity.

Using THEMIS visible spectrum image data for part of Phlegethon Catenae, we mapped pit crater outlines, surface drainage channels, and the footwall cutoff traces for all visible faults (Fig. 3A). East- and west-dipping faults are approximately evenly distributed, producing a system of horsts and grabens. Crosscutting relationships indicate that pit craters postdate fault scarps and drainage channels (e.g., Fig. 3A). Drainage channels radial to Alba Patera

are cut by faults with vertical offset and no consistent evidence of strike-slip displacement. In several cases, faults cut impact craters and there are no clear examples of impact craters postdating fault scarps. Pit craters can be observed at all stages of formation. One area where incipient pit crater formation can be observed is in Phlegethon Catenae (NW corner of Fig. 1B). Here, the smallest pits have apparently flat floors with surface textures similar to the surrounding topographic surface; the steeper pit walls are smooth. We interpret these as being

youthful pits where surface subsidence has not progressed so far as to have destroyed the original surface of in-falling material. This is supported by the location of these pits in relation to larger pits to the SW and NE. The overall geometry of this system suggests a breached relay structure between the propagating tips of the larger-displacement fault systems (e.g., Ferrill et al., 1999a).

Few if any impact craters are observed within pit craters (Fig. 1). Several large pits appear to have conical forms with no evidence of wall erosion or sediment accumulation. These observations suggest that the pit craters are among the youngest features on Mars. Eolian sediment transport continues to be active on Mars (Malin and Edgett, 2001); therefore, the lack of sediment accumulation and erosion in pit craters suggests that the pits have been active recently.

Although fault scarps appear to be steep, there is no reliable information to constrain the dip of the faults mapped. The consistent orientation of fault traces indicates that faulting took place in a uniform regional stress field. Slip-tendency analysis (Morris et al., 1996) of the Tractus Fossae fault system was performed applying a uniform normal-faulting effective stress tensor with an extension direction of 105° ($\sigma_1 =$ vertical; $\sigma_2 =$ azimuth 015° , 66% of σ_1 ; $\sigma_3 =$ azimuth 105° , 32% of σ_1). This stress tensor orientation produces the highest slip tendency along mapped faults with dips of 65° – 80° , would produce slip tendencies sufficient to overcome the frictional resistance to sliding on faults (Collettini and Sibson, 2001), and was the most probable orientation at the time of faulting (Figs. 3B and 3D). Dilation tendency analysis (Ferrill et al., 1999b) using the same stress tensor yields consistently high dilation tendencies for the same population of faults (Figs. 3C and 3E).

ANALOG MODELING

Based on our analysis of Mars data, we simulated slip on a normal fault with an upward steepening refracted cross-sectional profile beneath an undeformed and unconsolidated overburden. In these simulations, unconsolidated dry white or dyed sand represents Mars surficial materials (Sims et al., 2003). Constant thickness (1–3 cm) rigid wooden or aluminum plates, with or without an

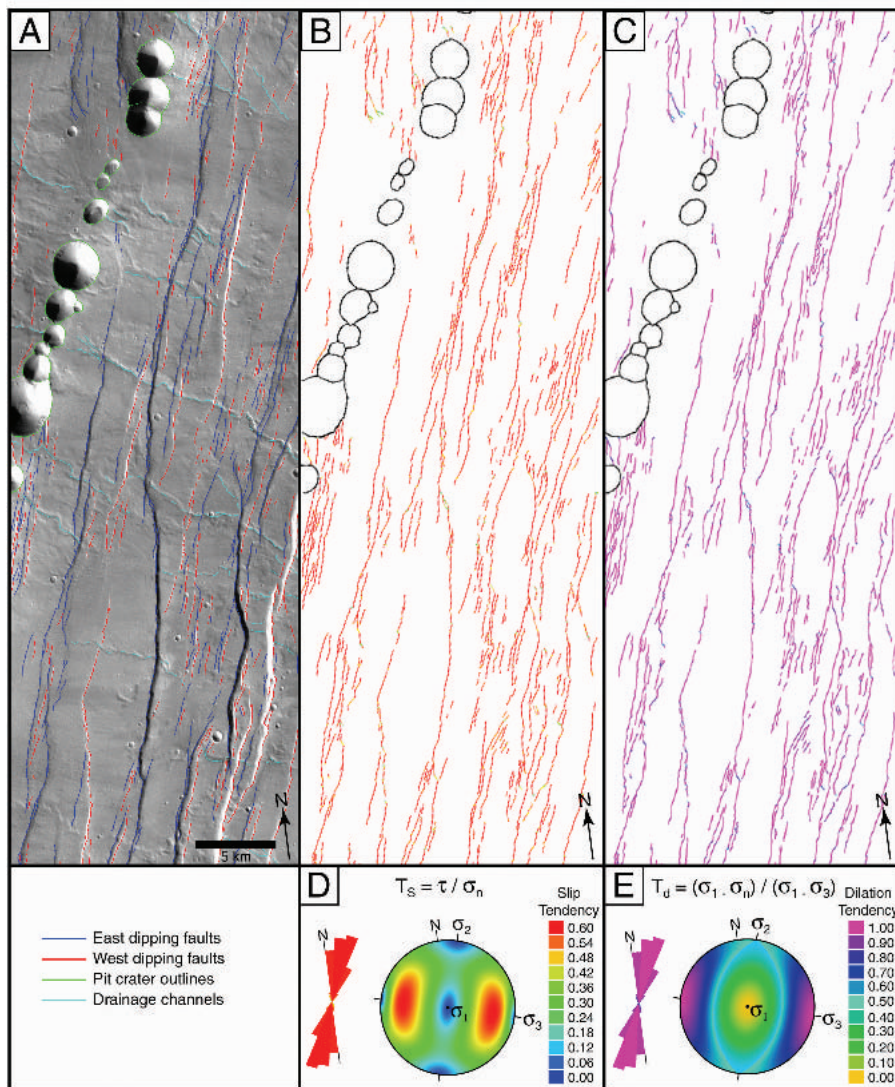


Figure 3. (A) Map of fault traces (footwall cutoff lines) in part of southeastern Alba Patera based on THEMIS visible spectrum image shown in Fig. 1. (B) Slip tendency (Ts) and (C) Dilation tendency (Td) maps of Alba Patera faults, illustrate the strong preferred orientations of normal faults in the mapping area. $\tau =$ maximum resolved shear stress, $\sigma_n =$ resolved normal stress, $\sigma_1 =$ maximum principal compressive stress, $\sigma_3 =$ minimum principal compressive stress. (D) and (E) contain lower hemisphere equal angle stereographic projection plots and length weighted rose diagrams for slip tendency analysis and dilation tendency analysis, respectively.

overlying layer of cohesive powder, were used to represent dilating fissures beneath the surficial material. Two fault geometries with imposed displacements were used to simulate fault geometry at depth: (a) a vertical upper fault segment linked to a horizontal detachment at depth (Sims et al., 2003), and (b) a vertical fault segment linked to a fault with

65° dip (consistent with slip tendency analysis) at depth (Fig. 4A). Plates were initially edge-to-edge, and tabular voids were created by progressively separating the plates. In some experiments, a layer of cohesive powder, capable of supporting steep to near vertical walls, was used in combination with rigid plates to represent the refracted fault geometries.

Fault traces were slightly irregular or corrugated (Ferrill et al., 1999a), similar to faults observed on Earth and Mars.

Most pit chain morphologies observed on Mars, including (i) chains of isolated circular or irregularly shaped pits, (ii) elongate pits, (iii) pit chains grading to troughs along trace, and (iv) pit chains along graben bounding faults (Wyrick et al., 2004), were reproduced in our physical models. Modeled pit chains follow a uniform developmental sequence from an alignment of circular pits of varying size, and variable (Sims et al., 2003) or regular (Horstman and Melosh, 1989) spacings, to elongate pits, to coalesced pits that form continuous troughs. Different stages of development were often present simultaneously along fault strike (e.g., Fig. 4H). Model pit chains tend to form somewhat irregular map traces, even where formed over parallel rigid plates. In some experiments, pit formation was preceded or accompanied by the formation of elongate grabens (Sims et al., 2003), with pit formation usually occurring along a single bounding fault.

In Model 19NOV03b (Fig. 4), displacement was gradually applied along the 65° fault segment at depth and produced oblique dilation of a vertical fault segment beneath a layer of cohesive powder. Overlying the cohesive powder was a layer of dry sand. As displacement progressed, an extension fracture propagated upward through the cohesive

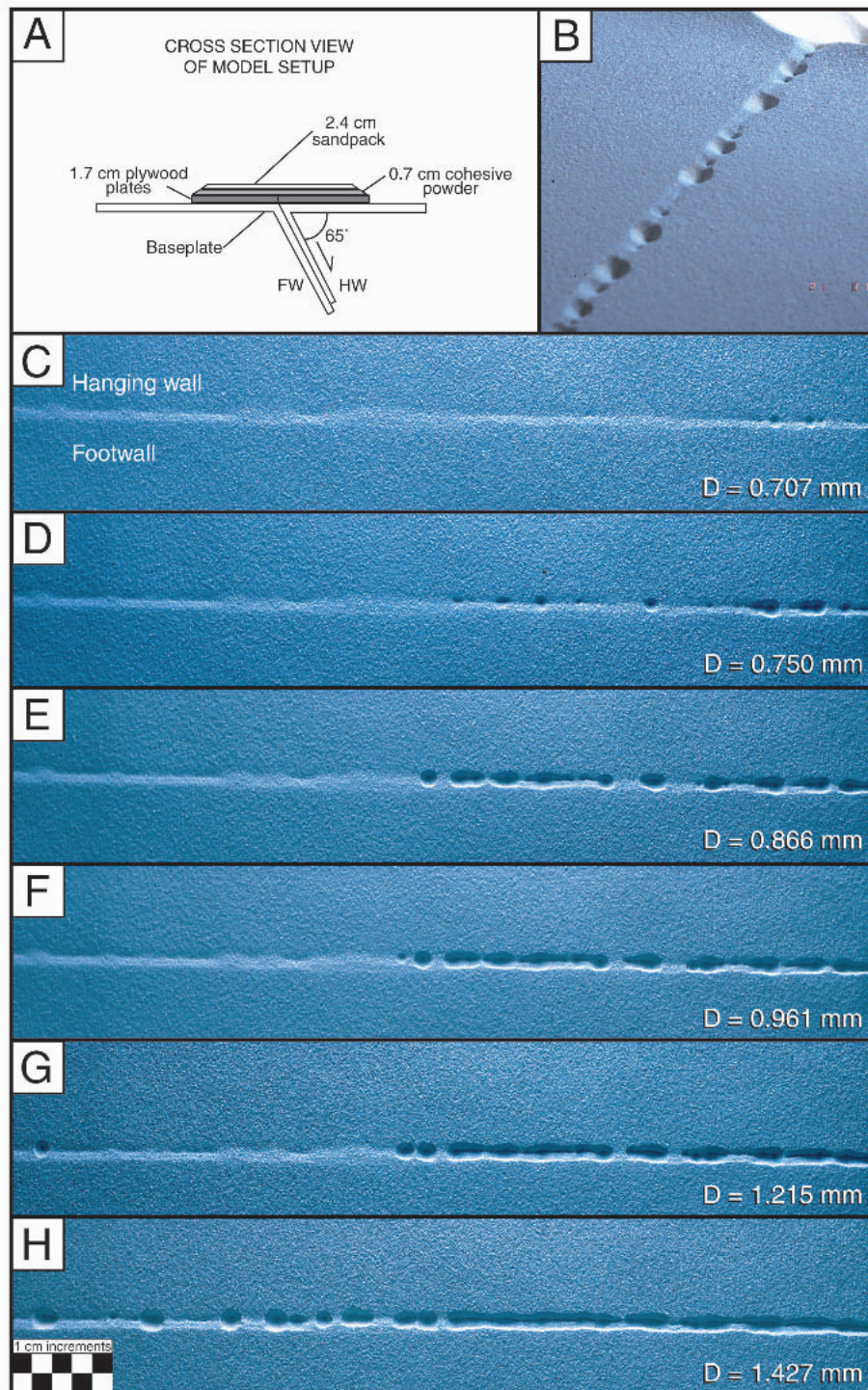


Figure 4. Analog modeling example of normal faulting and pit chain development. Model 19NOV03b produced a pit chain in unconsolidated sand above a refracted normal fault. (A) Schematic cross section illustrating model setup for simulation of pit chain development in a sandpack resting on a layer of relatively cohesive silica powder above a refracted normal fault that is vertical at the top and dips 65° at depth. Imposed displacement is normal dip slip along the 65° fault segment. The vertical fault segment experiences oblique dilation parallel to the slip vector on the 65° fault segment at depth. (B) Oblique photograph illustrating model pit chain development along the surface rupture trace of the normal fault. (C–H) Plan view photographs (illumination is from the top in each photo) of the model surface during progressive displacement, D, of the underlying normal fault.

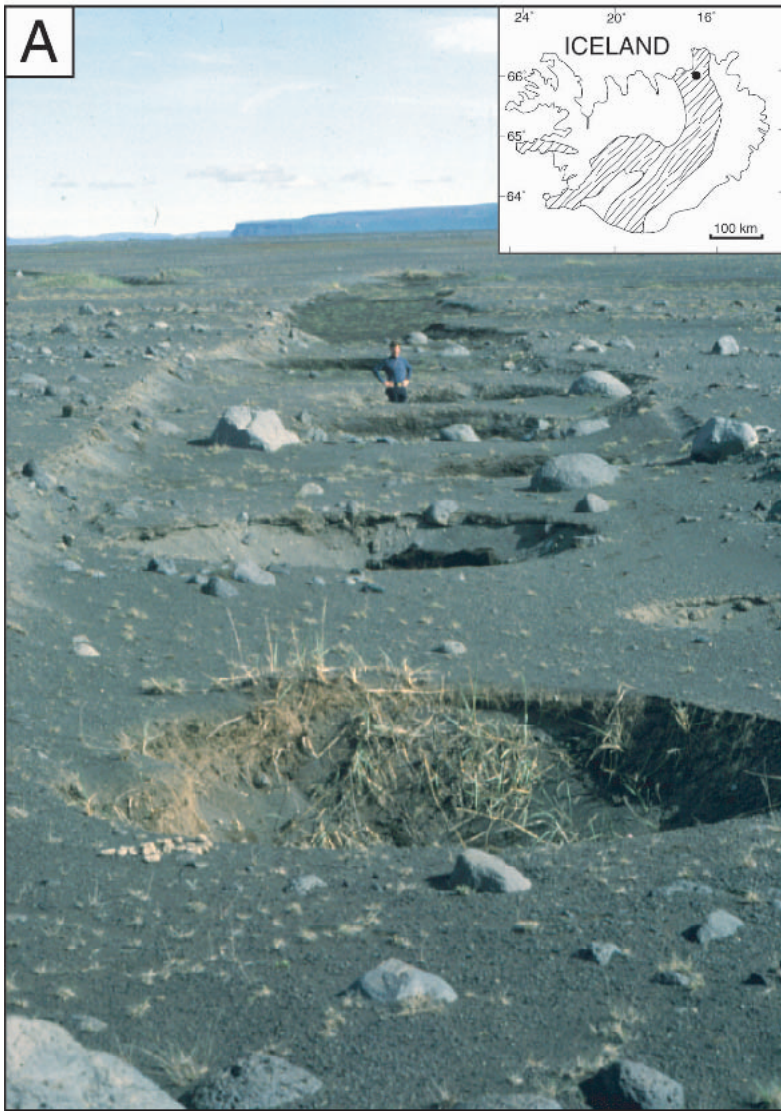


Figure 5 (here and on facing page). Pit chains in northeastern Iceland. (A) Shallow grabens and pit chain in gravels overlying basalt near the north coast of Iceland (inset map of Iceland shows the extent of the mid-Atlantic Ridge, and dot shows location of photograph near Ásbyrgi). Grabens and pit chains developed in response to an extensional deformation event in December 1975. Photograph was taken by George McGill on July 3, 1976. View is to the north. Note person (Bran Potter) standing in pit. (B) Aerial photograph taken on September 5, 1958, shows no evidence of pit chains and only subtle evidence of faulting. (C) Aerial photograph on September 2, 1976, shows pit chains. (D) Aerial photograph taken on August 3, 1984, shows growth and development of several pits between 1976 and 1984. Also during the 1976–1984 period, a stream channel established itself along one pit chain (trough), and subsequent erosion and deposition progressively erased the surface expression of the pit chain along the stream’s course.

powder layer and progressively dilated. The void space filled by draining of the overlying unconsolidated dry sand (Fig. 4), producing pits up to 1 cm diameter in the sand. Initially isolated pits gradually grew and intersected, forming elongate pits, scalloped sided coalesced pits, and relatively smooth-sided troughs (Figs. 4C–4H).

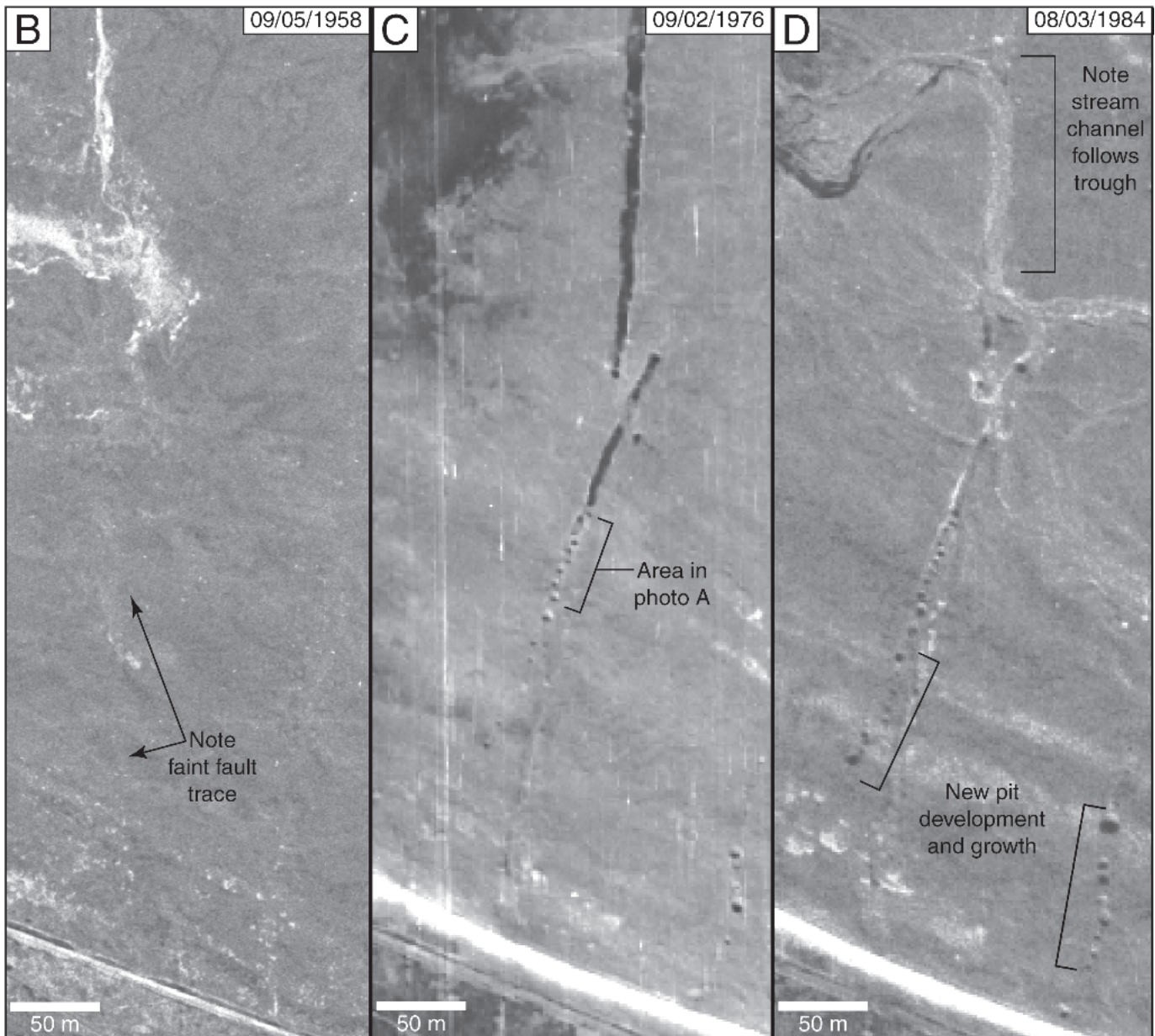
PIT CRATER CHAINS IN ICELAND

Although pit crater chains are common on Mars, comparable examples on Earth are rare (Wyrick et al., 2004). A striking example of pit chain development occurred in Iceland during the rifting episode that started in 1975. December 20, 1975, marked the beginning of a rifting and subsidence episode along the Mid-Atlantic spreading center

in northeastern Iceland. This first event, the largest event of the deformation episode, continued into February 1976 and was followed by 10 more events from September 1976 to May 1979 (Sigurdsson, 1980). Deformation events were marked by seismicity, small-volume volcanic eruptions at Krafla (during first, fifth, and sixth events), reactivation of faults along existing surface traces, grabens, and pit crater chains in alluvium where surface ruptures were previously unmapped. Seismicity and ground surface fissures associated with the December 1975 through February 1976 event were concentrated in a NNE-SSW-trending elongate region extending 30–50 km north of the Krafla caldera center. Surface ruptures with components of

dilation and normal dip slip (Opheim and Gudmundsson, 1989; Angelier et al., 1997; Dauteuil et al., 2001) formed along previously mapped fault traces in basalt flows. To the north, a 1–1.25-km wide and 12-km-long system of fault ruptures, fissures, and pit craters also formed in the floodplain deposits of Jökulsá á Fjöllum (Sigurdsson, 1980). In alluvium of the Jökulsá plain, within 2 km north of the contact with the gently northward-dipping basalt flows, the normal fault system reactivated by the first event produced a 1-km-wide graben that accommodated an estimated 1.5 m of horizontal extension and approximately 1 m vertical throw. The western graben-bounding fault system accommodated approximately 30 cm horizontal extension and 1 m vertical throw.

George McGill provided us with the ground photograph of a pit crater chain formed within this western graben-bounding fault system in the floodplain deposits of Jökulsá á Fjöllum during the December 1975 through February 1976 event (Fig. 5A). Aerial photographs taken on (1) September 5, 1958, (2) September 2, 1976, and (3) August 3, 1984, of the same region document the formation and evolution of this pit crater chain (Fig. 5). The photograph taken on September 5, 1958, (Fig. 5B) shows no evidence of pit chains and only subtle evidence of faulting. The aerial photograph taken on September 2, 1976,



(Fig. 5C) shows pit chains that trend NNE-SSW and consist of pits ranging in size from ~1 m to ~10 m in diameter. Some pit chains are clearly associated with normal faults, in some cases occurring within a down-dropped graben between two normal fault scarps. Aerial photographs taken on August 3, 1984, show growth of numerous pits between 1976 and 1984. Also during the 1976–1984 period, a stream channel established itself along one pit chain (trough), and subsequent erosion and deposition progressively erased the surface expression of the pit chain along the stream's course. Aerial photographs taken after August 3, 1984, show some growth of

pits and minor change in the amount of erasure of pits by the stream erosion and deposition.

Although the pit crater chains formed during this crustal deformation episode in Iceland are small, generally <10 m diameter, they are excellent geologic analogs for pit craters observed on Mars. Faults along the Icelandic rift have dilational displacement (Opheim and Gudmundsson, 1989; Angelier et al., 1997; Dauteuil et al., 2001) and this is likely the result of the contrast in strength between basalt and interbedded ash and sediment (Ferrill and Morris, 2003). Where dilational jogs occur beneath unconsolidated surface material,

they cause subsidence of this material into fault-generated void space and hence pit crater formation. We interpret their small size to be due to the thin layer of unconsolidated surficial deposits overlying the basalt and a small void volume formed by 10 cm to 2 m fault displacements. In Iceland, pit crater volumes represent a short series of events and rapid sedimentation fills the pit crater before the next deformation event (Fig. 5D).

WHY ARE PIT CRATERS SO LARGE AND SO EVIDENT ON MARS?

On Mars, where erosion is minimal, pit size may represent a cumulative

volume as a result of many events spread out over time. On Earth, erosional and depositional processes may erase and/or subdue the surface expression of pit craters within decadal time scales. Although surface processes such as wind and meteorite impacts can erase pits, these processes are less effective than flowing water, which is the dominant erosive agent on Earth. Aerial photography from Iceland (Fig. 5) demonstrates the important role of erosion and deposition in the obliteration of terrestrial pit craters. Perhaps more important, the influence of Martian gravity on faulting in the brittle crust plays a major role in the overall geometry of Martian faults. Martian gravity at 3.72 m/s^2 is substantially less than Earth's gravity at 9.81 m/s^2 (Table 1 in Esposito et al., 1992) and this influence, along with rock density, controls the lithostatic stress profile, which is the primary influence on failure characteristics of rock at depth.

FAILURE MODES, NORMAL FAULT PROFILES, AND GRAVITY

Active normal faults on Earth commonly have steep dips near the surface and progressively more gentle dips at depth (Walsh and Watterson 1988). Dips of 70° – 90° are common in the uppermost brittle crust (0–2 km) in a variety of rocks including volcanic tuffs (Day et al., 1998), basalts (Opheim and Gudmundsson, 1989; Angelier et al., 1997; Dauteuil et al., 2001), limestones

(Ferrill and Morris, 2003; Ferrill et al., 2004), and clastic sedimentary rocks (Walsh and Watterson, 1988). Dips of 60° are common at depths of 2–5 km, and dips of 35° – 55° are common in the lower part of the brittle crust ($>5 \text{ km}$; Jackson and White, 1989; Collettini and Sibson, 2001). Dips of 0° – 35° are associated with detachments at the brittle-ductile transition in the crust and in extremely weak sedimentary layers.

Fault orientation in rock is initially controlled by mechanical properties of the faulted rocks, and magnitudes and orientations of the effective principal stresses at the time of failure. Failure angle is the angle between the failure plane and the maximum principal compressive stress at the time of failure. The angle of shear failure (θ) in rock can be calculated from the angle of internal friction (friction angle, ϕ) using:

$$\theta = \pm (45^\circ - \phi/2) \quad (1)$$

Failure mode depends on the differential stress ($\Delta\sigma = \sigma_1 - \sigma_3$; where $\sigma_1 =$ maximum principal compressive stress, and $\sigma_3 =$ minimum principal compressive stress) and the effective minimum principal compressive stress (σ_3' ; $\sigma_3' = \sigma_3 - P_f$, where $P_f =$ pore fluid pressure) at the time of failure, and upon the strength characteristics of the rock (e.g., Mandl, 1988 and references therein; Fig. 6A). There are three types of failure (Mandl, 1988; Ferrill and Morris, 2003). Shear

failure is characterized by a failure angle given by Equation (1), displacement parallel to the fracture surface, and the effective normal stress (σ_n') acting on the fracture surface, $\sigma_3' < \sigma_n' \geq 0$ (Fig. 6A). Tensile failure is characterized by a failure angle of 0° and displacement perpendicular to the fracture surface (Fig. 6A). In this case $\sigma_3' = \sigma_n' < 0$ and $|\sigma_3'| \geq$ tensile strength of the rock. Hybrid failure is characterized by a failure angle between 0° and the angle calculated by Equation (1), and displacement oblique to the fracture surface ("dilatant faults," Mandl, 1988; Fig. 6A). Hybrid fracture, hypothesized and debated for decades, has recently been demonstrated in laboratory experiments (Ramsey and Chester, 2004). In the case of hybrid failure, $\sigma_3' < \sigma_n' < 0$ and $|\sigma_3'| <$ tensile strength of the rock.

The normal faulting regime is defined by a vertical maximum principal compressive stress (Anderson, 1951). The vertical or lithostatic stress (σ_v) increases with depth and is a function of the thickness (h) of the overburden, the integrated density of the overburden (ρ), and the acceleration due to gravity (g):

$$\sigma_v = \rho gh \quad (2)$$

As a result of the vertical stress gradient, differential stress required for failure also increases with depth. Assuming a range of density values of 2.7 – 3.1 g/cm^3 for basalt (Carmichael, 1989) as examples,

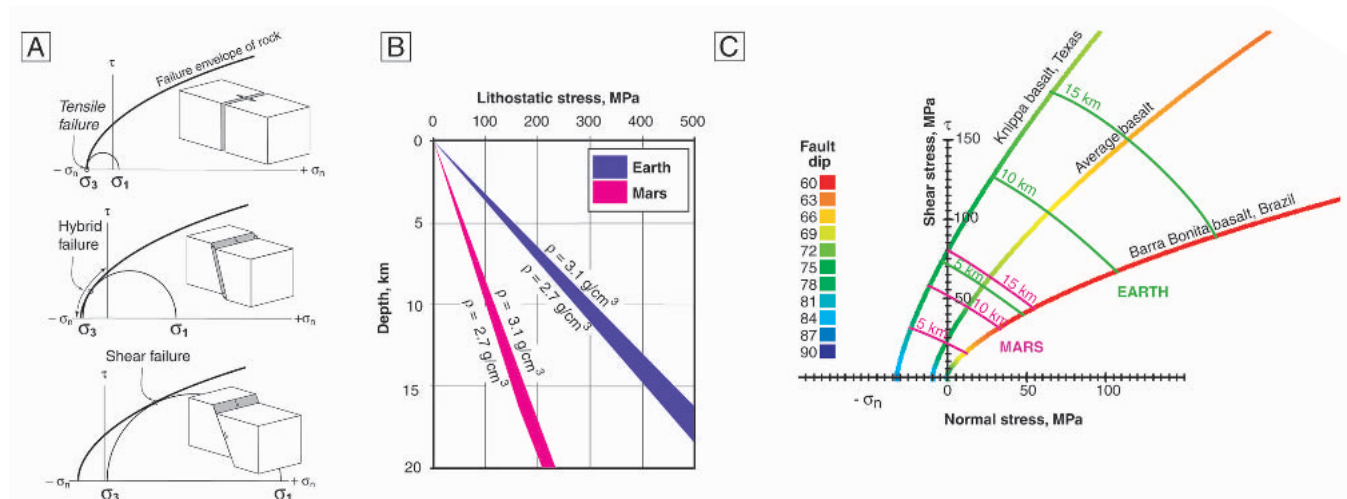


Figure 6. (A) Mohr circle and illustration of tensile, hybrid, shear failure modes (after Ferrill and Morris, 2003). (B) Lithostatic stress gradients on Earth vs. Mars assuming same rock type (basalts) and no hydrostatic gradient. (C) Hoek-Brown failure envelopes for basalt, color coded according to fault dips for the normal faulting regime, comparing depths of occurrence in Earth (green) versus Mars (red).

lithostatic stresses can be calculated as a function of depth for Earth ($g = 9.81 \text{ m/s}^2$), and Mars ($g = 3.72 \text{ m/s}^2$) (Fig. 6B). Hoek-Brown failure envelopes (Hoek and Brown, 1988; Schultz and Zuber, 1994) were constructed for an average basalt and two end-member basalts based on failure test data from 37 basalts (Gevantman, 1982). Basalt was chosen as a likely rock type on Mars because extensive volcanic plains cover half of its surface (Gornitz, 1997). Each point on a failure envelope represents a unique value for $\Delta\sigma$, and because the envelopes are nonlinear (especially at low values of $\Delta\sigma$), each point also represents a unique value of the failure angle. Because $\Delta\sigma$ for failure is depth-related, and ignoring the effects of pore fluid pressure, the envelope curves can be color-coded for normal fault dip (normal fault dip = $90^\circ - \text{failure angle}$) and correlated with depth in the crust for both Mars and Earth (Fig. 6C). Tensile and hybrid failure for an "average basalt" and corresponding fault dips of 78° – 90° , would be limited to $<2 \text{ km}$ on Earth, but are likely to depths of $\sim 5 \text{ km}$ on Mars.

Fault profiles on Mars are likely to be similar in shape to analogs on Earth, but depth-dependent fault dip transitions are likely to occur deeper than on Earth and occur more gradually with depth. The seismogenic depth associated with extensional deformation on Mars is expected to be shifted deeper due to lower gravity, but would also depend on geothermal gradient at the time of active faulting. Dilational faulting, found in the upper $\sim 2 \text{ km}$ on Earth, may extend to depths of 5 km or more on Mars due to the factor of 2.64 difference in gravity.

DISCUSSION

The common occurrence of pit crater chains on Mars compared with Earth is probably controlled by two factors: weathering and erosion rates and faulting style. Earth's surface is reworked much more rapidly than that of Mars because of the thicker terrestrial atmosphere and the prevalence of surface and near surface water. Pit crater chains formed in Iceland have been significantly modified and even obliterated within a decade of formation. Martian pit craters are larger by at least an order of magnitude than any on Earth. This size difference has three likely sources: (i)

thickness of low-cohesion surface material, (ii) depth-extent of the dilational fault segment generating the subsurface void, and (iii) Martian pit craters represent a cumulative volume over a number of faulting events, whereas on Earth, pit crater volume represents a single event. Normal faults and extension fractures on Mars are more likely to be dilatant and extend to greater depths than similar features on Earth because of the lower Martian gravity and consequent lower stress gradient.

Martian thermal history may also have influenced the prevalence of dilational faults. Mars has cooled and consequently the thickness of the Martian cryosphere and the thickness of ice in the cryosphere has increased over time (Clifford and Parker, 2001). At a latitude of 30 – 40°N on Mars today (e.g., Alba Patera examples in Fig. 1), the top of ice is estimated to be at a depth of 10 m (Fanale et al., 1986) and extend to a depth of 3.5 – 5 km (Clifford and Parker, 2001). The strength of unconsolidated material increases as freezing occurs (p. 272 in Sowers, 1979), therefore steepening the failure envelope. Consequently, the shear failure angle decreases and the stress range over which hybrid failure can occur increases, resulting in steeper dips for faults forming in a frozen stratigraphic section. Deeply frozen rock and unconsolidated materials (Carr, 1996; Clifford and Parker, 2001) would generate deeper, steeper dilatant fault segments. This effect on faulting would be more pronounced as the planet cooled.

Faults and fractures form important pathways for and barriers to groundwater flow on Earth (e.g., Ferrill et al., 2004, and references therein). We anticipate a similar relationship between these features and past groundwater flow on Mars. The occurrence of dilational faulting on Mars indicates highly permeable zones and potentially important sites for storage of water or ice and locations of mineralization. Pit chains indicate the presence (either now or formerly) of subsurface voids and potential conduits for groundwater flow, whether unfilled or filled with porous material. Understanding their mechanism(s) of formation is essential for constraining volumes and locations (depth and lateral extent) of groundwater flow conduits and potential areas of water storage,

with major implications for the subsurface distribution and movement of water and mineralization in the Martian crust. Recent description of outflow channels sourced by pits supports the interpretation of pit chains as indicators of important groundwater flow and discharge pathways (Dinwiddie et al., 2004). This interpretation is also consistent with outflow channels sourced by fractures (see summary in table IV of Clifford and Parker, 2001).

CONCLUSIONS

Many pit crater chains on Mars are produced by dilational normal faulting. Pit craters are larger and better preserved on Mars than on Earth because of more active erosion and deposition on Earth and the lower gravity and greater vertical extent of the zone of dilational faulting on Mars. In many areas, pit crater chains appear to be some of the youngest features, postdating drainage channels, faulting, and impact craters. Based on crosscutting relationships, pristine pit morphologies, and lack of evidence of sediment accumulation in pits, we believe that some pit craters may be actively forming. Dilational faults, whether partially filled or open cavities, are likely to have served as conduits for past groundwater flow and would now serve as reservoirs for water or ice.

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